



06/June/2019

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KISS

A spectrum-imager dedicated to the study of the secondary anisotropies of the CMB



Outlines

The path

- 
- The background of the slide is a photograph of the Teide volcano in Tenerife, Spain, viewed from the Teide observatory. The volcano is a large, conical mountain with a snow-capped peak, partially obscured by a green semi-transparent box containing the text. The foreground shows a rugged, rocky landscape with sparse vegetation under a clear blue sky with some light clouds.
- I. Science context and requirements**
 - Sunyaev Zel'dovich Effect
 - II. Observation strategy and instrument design**
 - Fourier Transform Spectrometry
 - Fast detectors: Kinetic Inductance Detectors
 - III. Laboratory tests and characterization**
 - Detectors performances
 - Geometrical characterization
 - IV. Installation and observations**
 - The Moon observations
 - V. Conclusions and perspectives**

Teide vulcano from Teide observatory

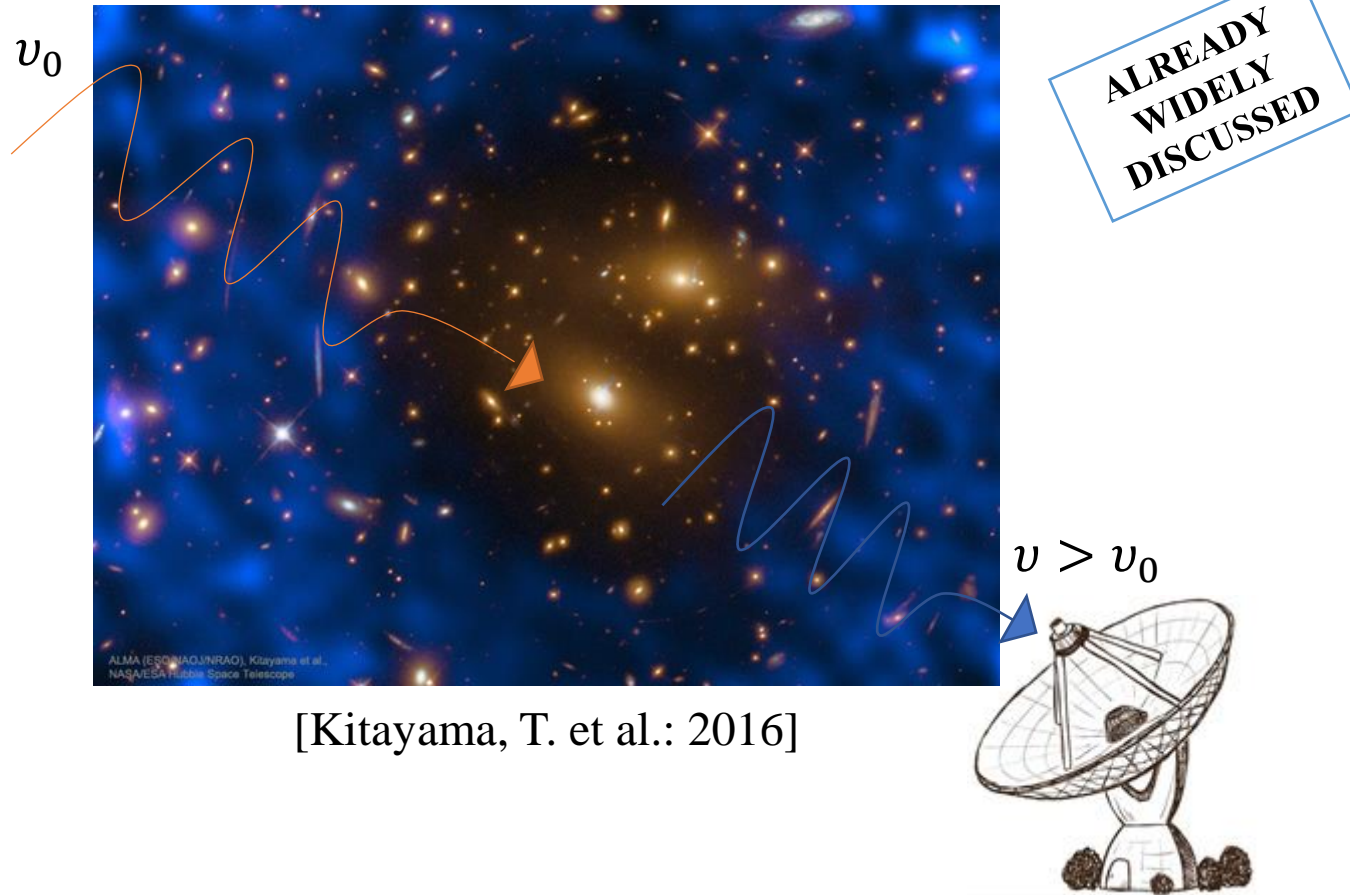
I. Science context and requirements

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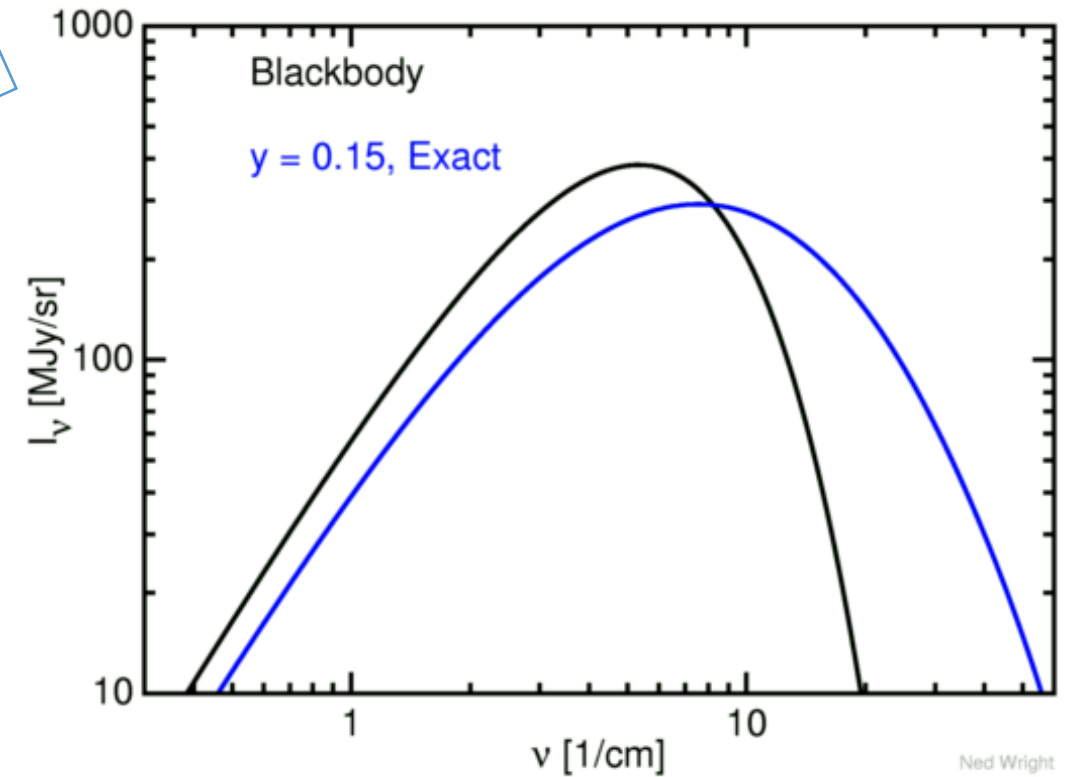
Sunyaev Zel'dovic Effect

Inverse Compton scattering CMB-IntraClusterMedium

RXJ1347.5



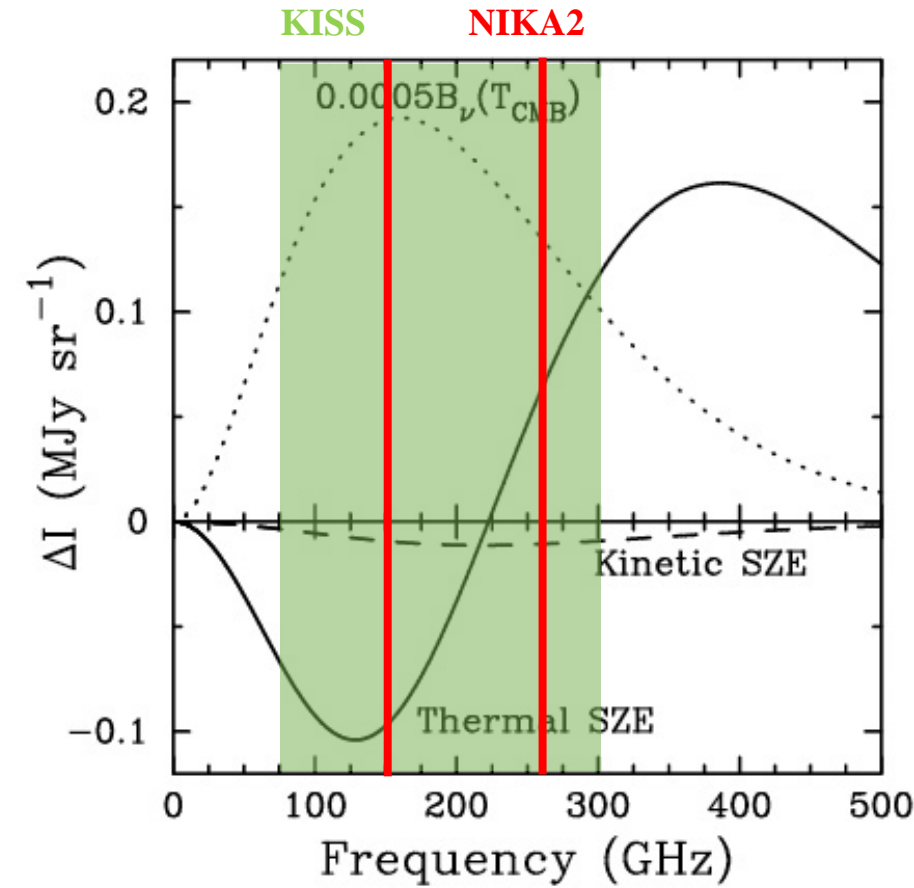
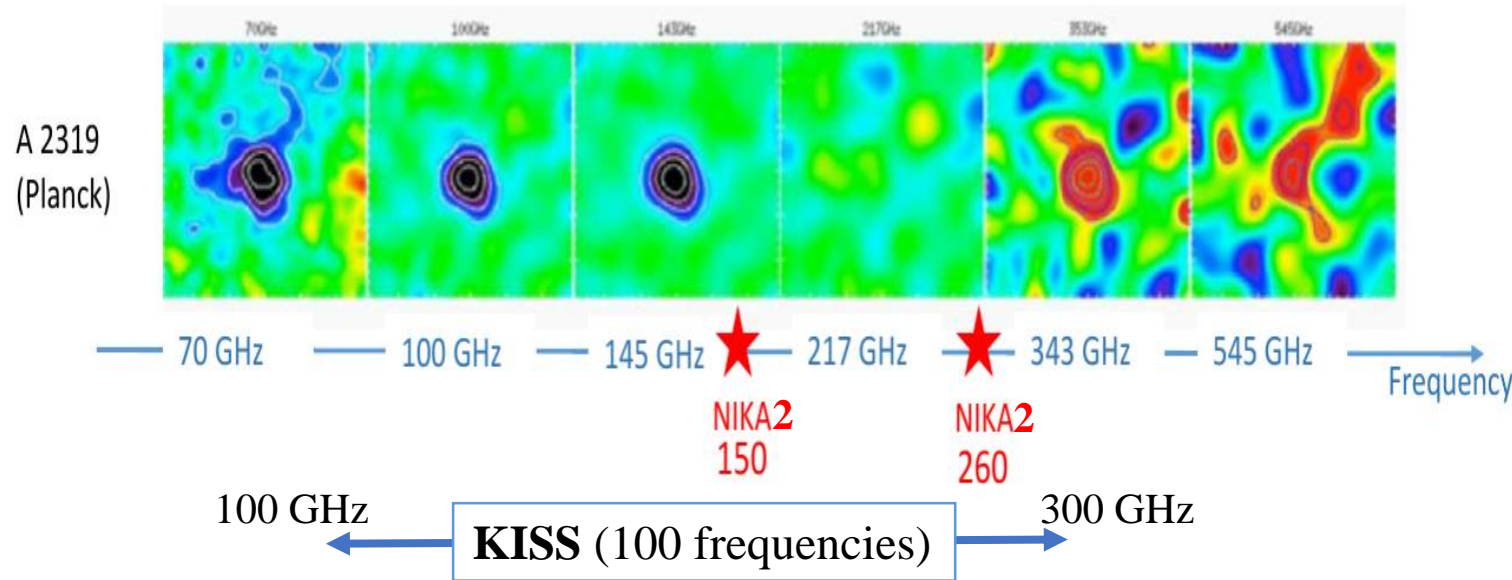
**Main target of
the experiment**



[Wright, E., L.: 2007]

I. Science context and requirements

SZE - state of art



Experiment	# frequencies	angular resolution @ 150 GHz	# pixels
Planck	9	5 arcmin	52
NIKA2	2	20 arcsec	3'000
KISS	100	4 arcmin	600
CONCERTO	150	50 arcsec	3'000

Observing several frequencies to separate the SZ components

I. Science context and requirements

State of art and necessities

GOAL

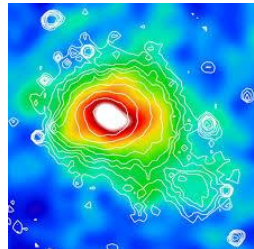
Low resolution spectroscopy observations of known **low redshift galaxies** at **mm wavelengths** to map cluster physical properties from spectral distortions.

STRATEGY

Compensate relative expected low sensitivity with respect to Planck or photometric ground-based instrument by integrating longer (tens of hours per cluster).

Use spectroscopy to fully separate different components and extract physical information from spectral distortions: pressure, temperature, density, mass, LOS velocity

Possible target
COMA cluster



II. Observation strategy and instrument design

II. Observation strategy and instrument design

From scientific requirements to instrumental characteristics

1) Low angular Resolution

low redshift clusters

2) Large FoV and band 100-300 GHz

~1 degree

3) Low Spectral resolution

~1.5-10 GHz at least 20 bins to separate properly different contributions

4) Maximum Sensitivity

photon noise detectors



1) Telescope : 2.5 m - Quijote

~ few degrees corrected FP angular resolution from about 2 to 5 arcmin

2) FTS Technique - Fast MPI

10 cm excursion , fast acquisition, avoid 1/f noise from the atmosphere (it is ground based)

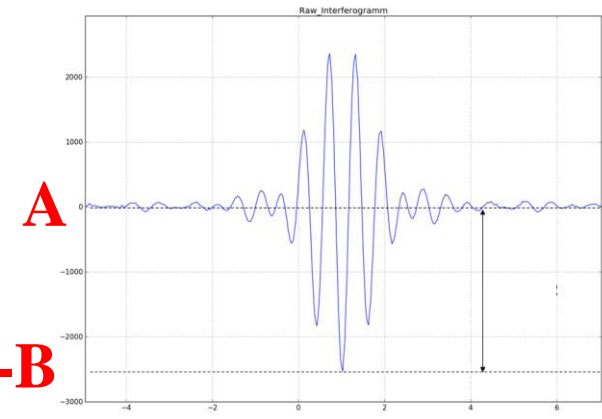
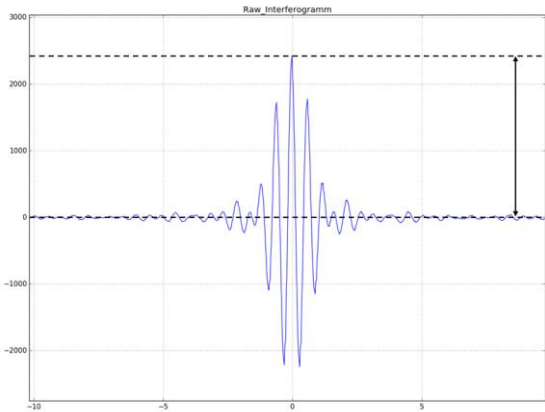
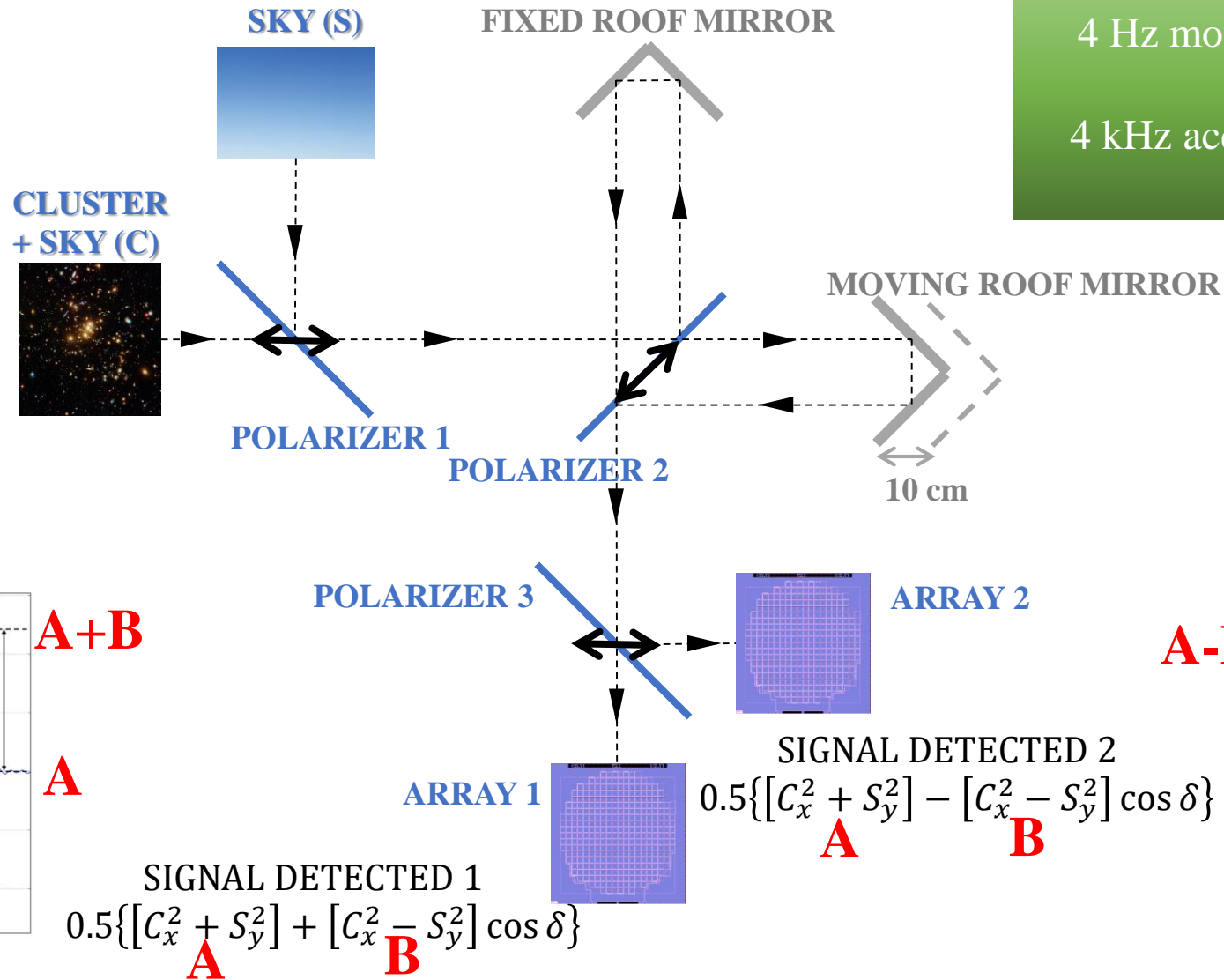
3) 2 Arrays of 300 pixels

II. Observation strategy and instrument design

Fourier Transformation Spectrometry

MPI definition:
a **Martin-Puplett Interferometer** measures the difference between the powers of the two input beams.

Necessity to be fast:
4 Hz moving mirror -> atmosphere fluctuations
4 kHz acquisition rate -> 100 points spectrum

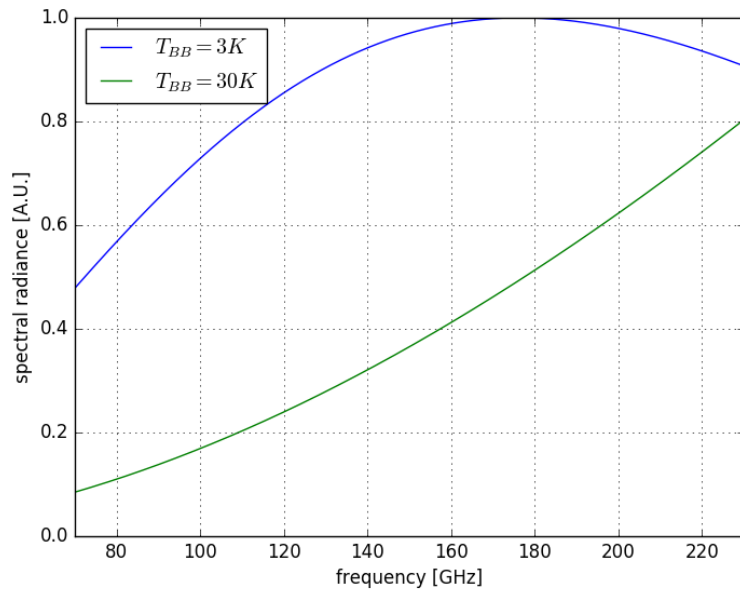


II. Observation strategy and instrument design

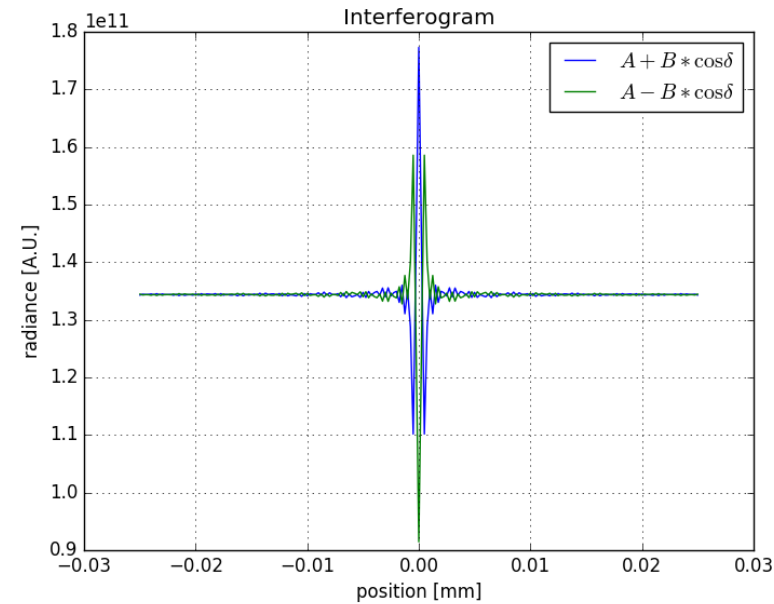
Fourier Transformation Spectrometry

e.g.: naïf simulation of the MPI

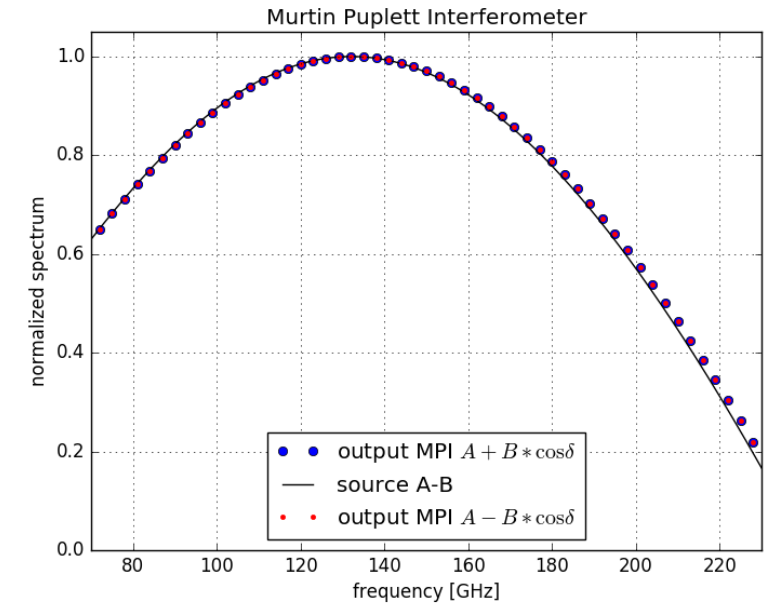
INPUT SIGNAL



INTERFEROGRAM



SPECTRUM



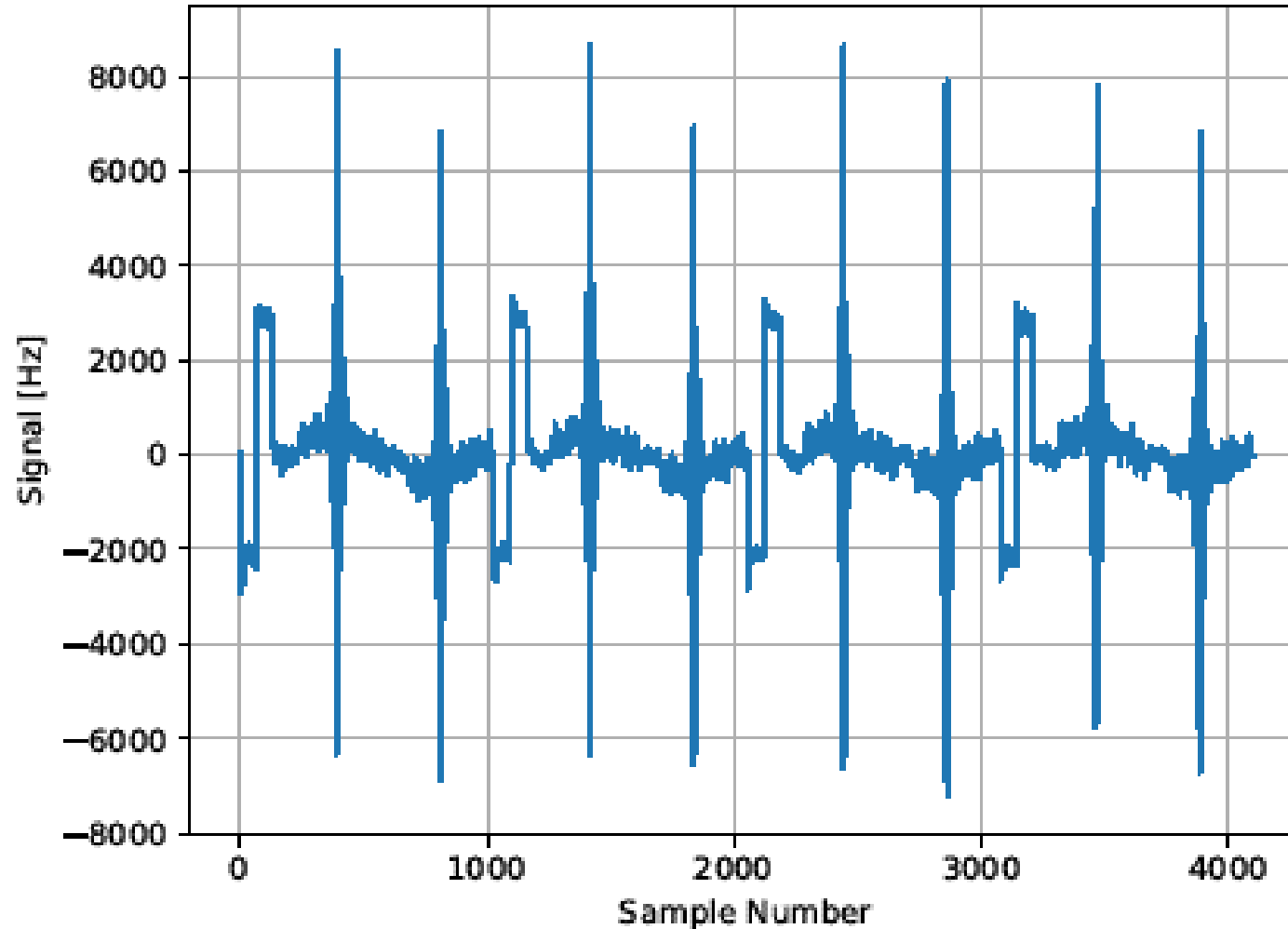
Interferometer

Fourier Transform

II. Observation strategy and instrument design

Readout technique: the tuning procedure

Real raw data – timeline modulation



**“Smart” solution
against
data inflation**

data rate:

NIKA2 ~ 1 Mb / s

KISS ~ 10 Mb / s

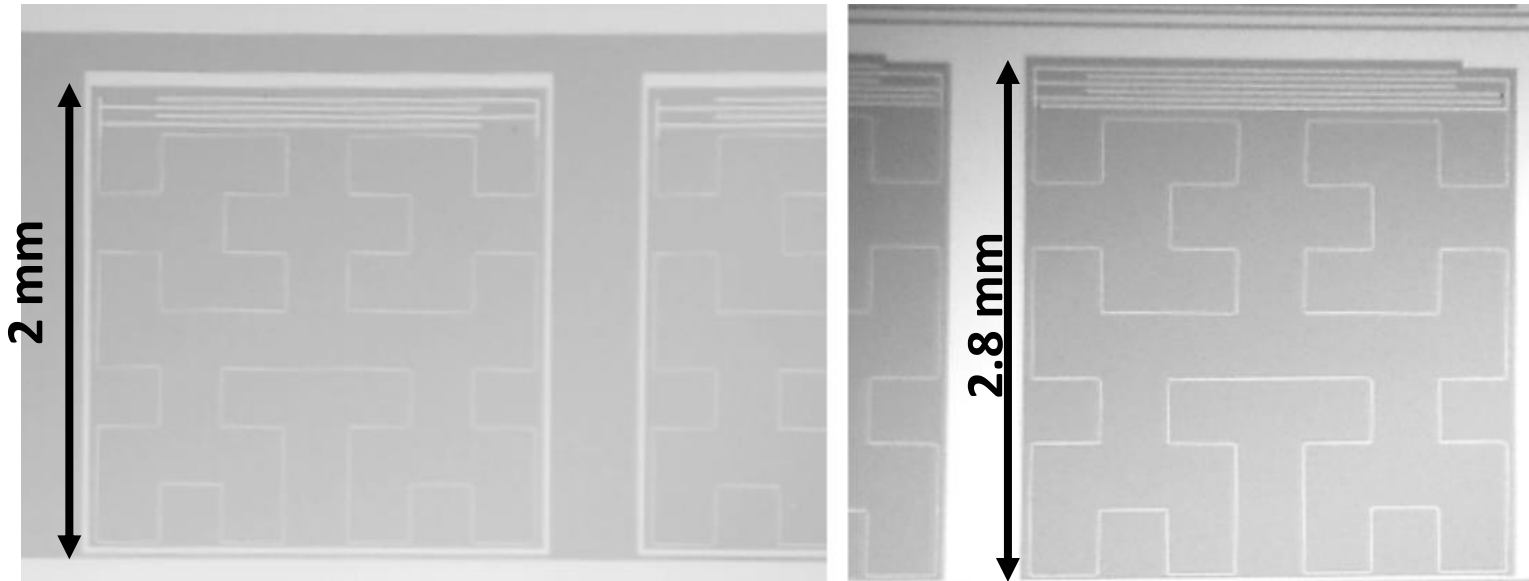
II. Observation strategy and instrument design

Fast Detectors: Kinetic Inductance Detectors

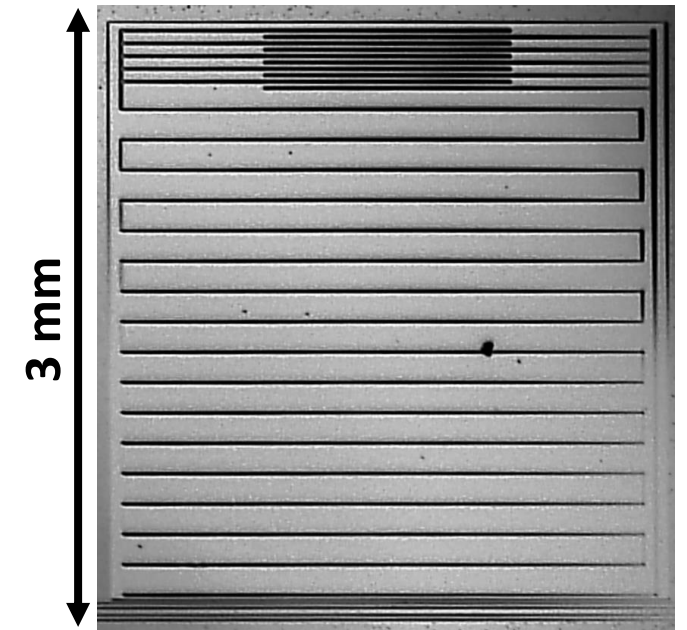
Time constant

$$\tau_{qp} \sim \text{tens } \mu\text{s}$$

NIKA2 (260 & 150 GHz)



KISS (70÷300 GHz)



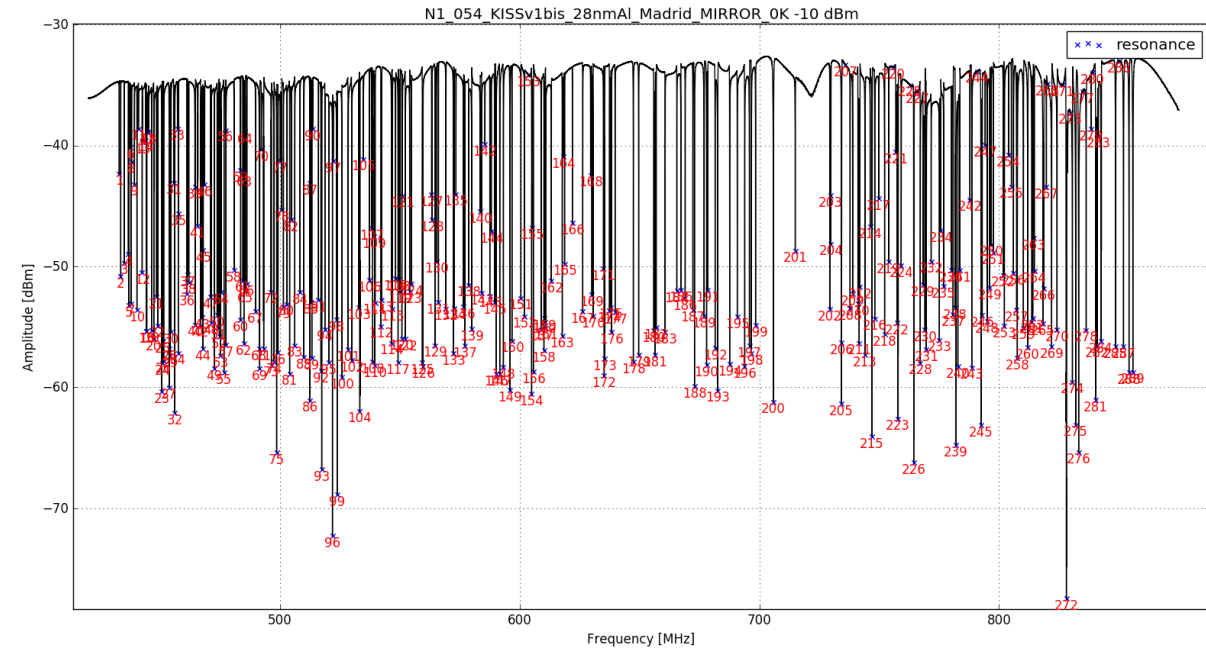
as seen in J. Goupy yesterday

Hilbert geometry:
not polarized, 260 GHz
received polarized light

Standard meander:
polarized, aligned with
MPI polarizer

II. Observation strategy and instrument design

Fast Detectors: Kinetic Inductance Detectors

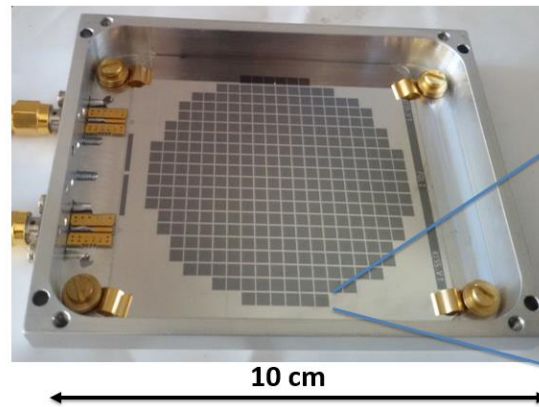


Multiplexed array
300 pixels

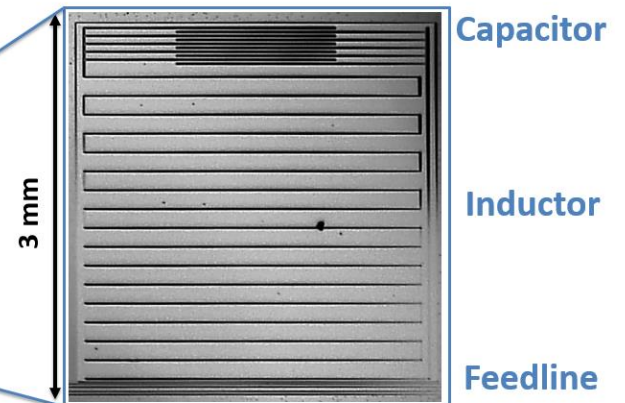
The realization of an array with a zoom over a pixel

KIDs reach the intrinsic limitation of sensitivity

Superconducting LC circuit with high Quality factor



Single KISS array of 316 pixels based on Ti-Al bilayers ($T_C \approx 0.95$ K).



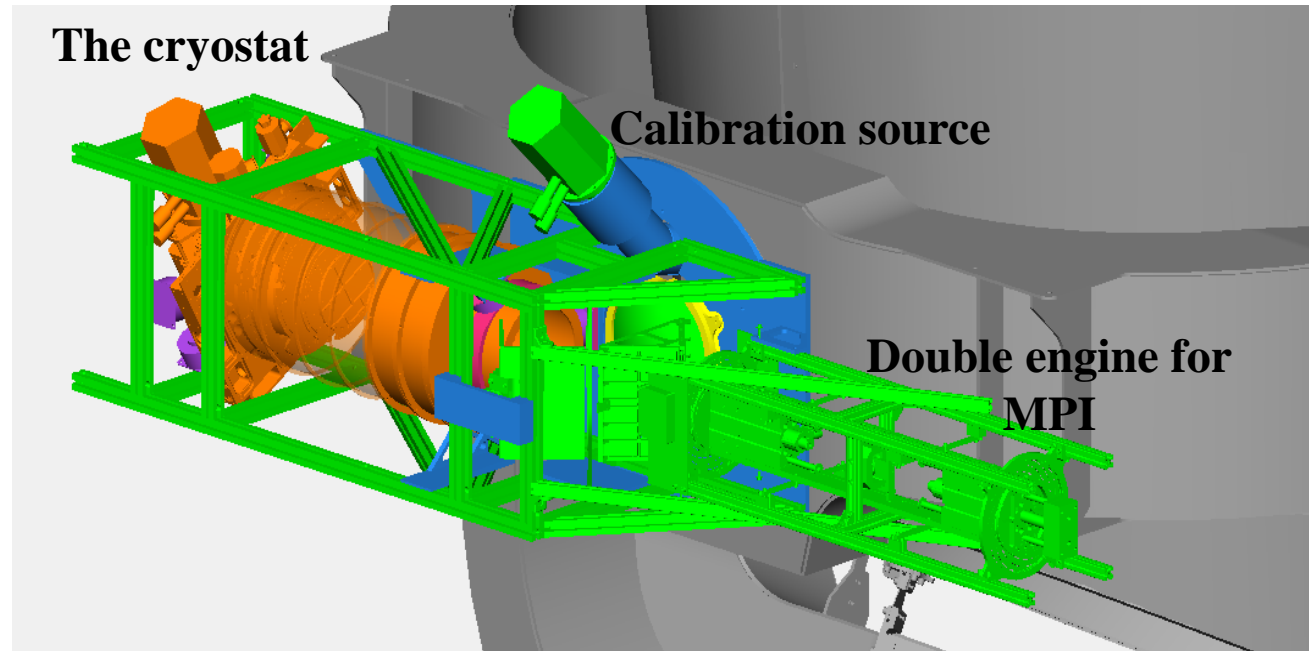
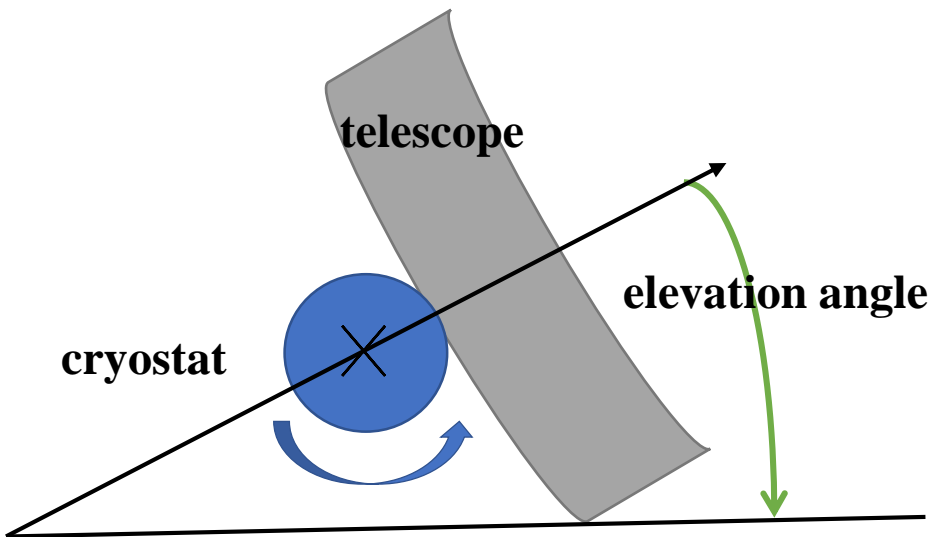
One of the 632 pixels within the KISS camera

II. Observation strategy and instrument design

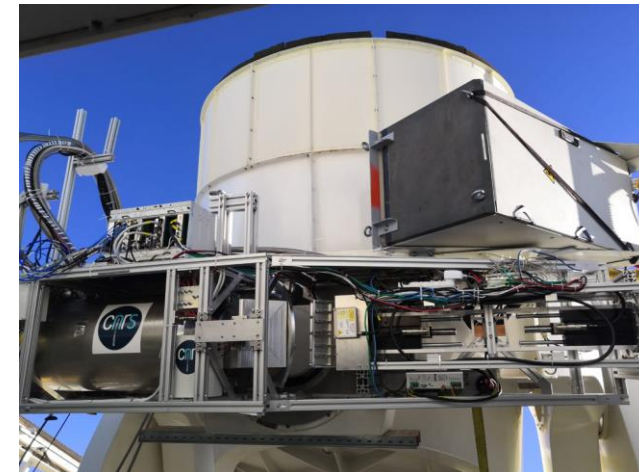
Fast Detectors: Kinetic Inductance Detectors

Technological constraints:

- calibration source for the MPI
- double MPI engine to delete the vibrations
- stable temperature cryostat for the focal plane



KISS installed at QUIJOTE telescope

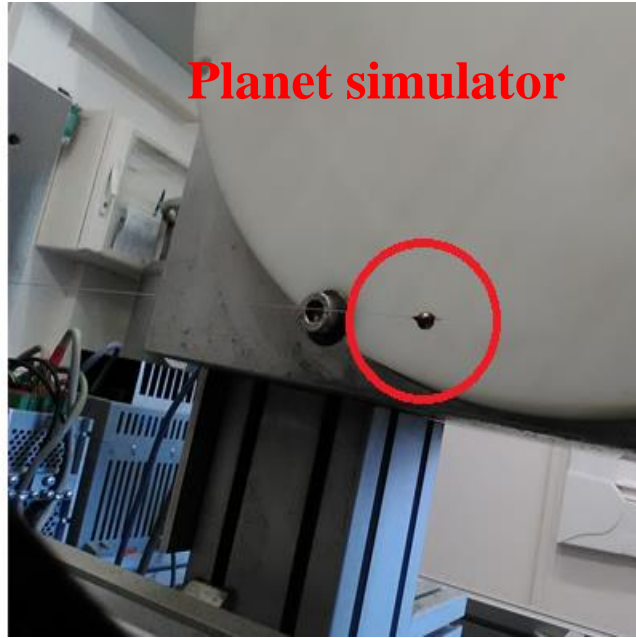


III. Laboratory tests and characterization

October 2017 – October 2018

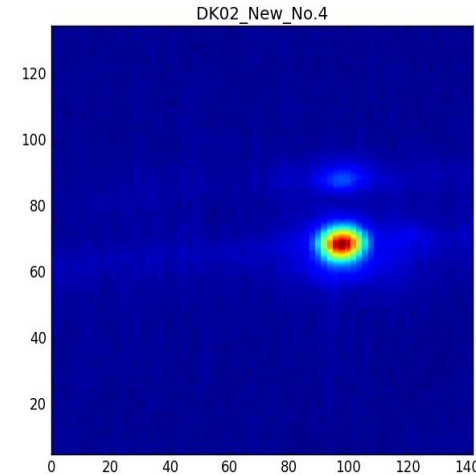
III. Laboratory tests and characterization

Geometrical characterization

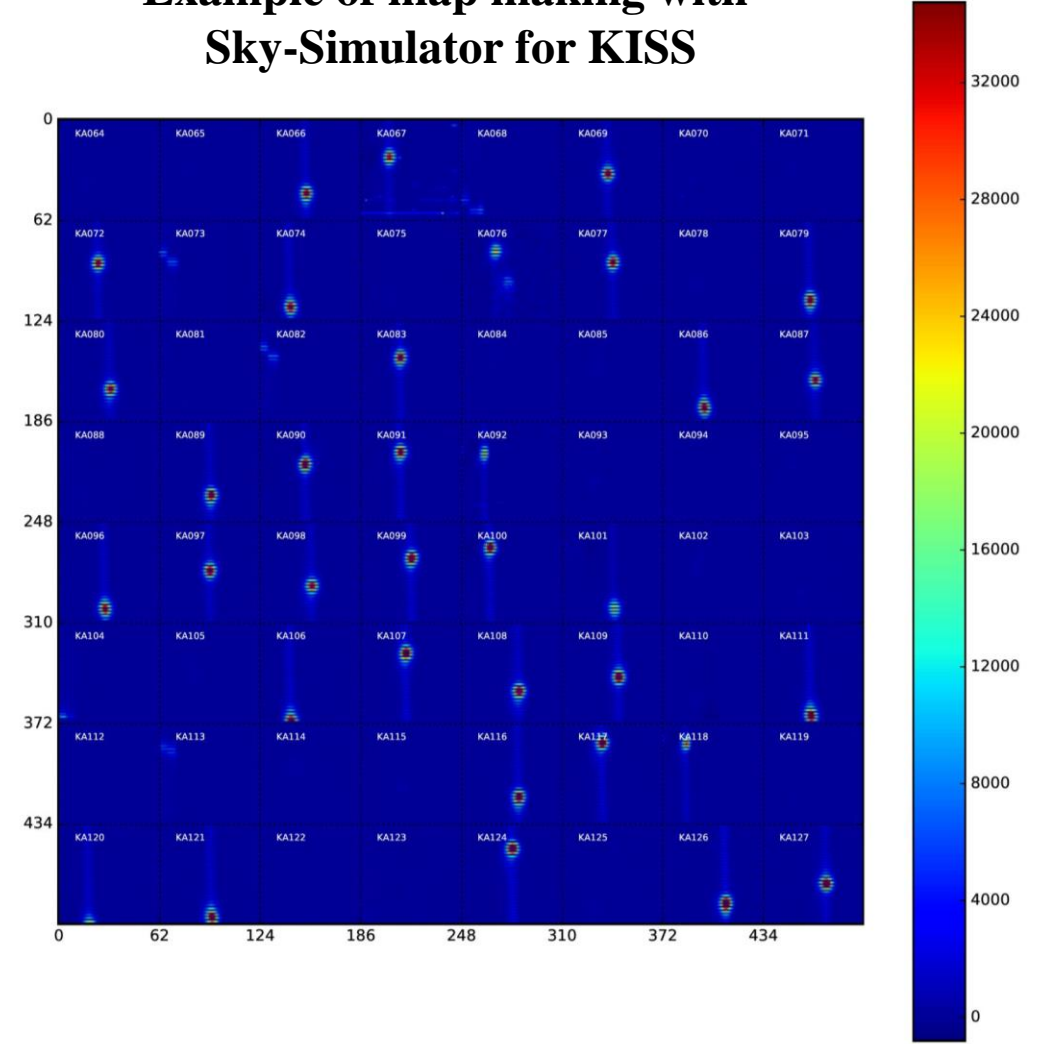


as seen in S. Shu yesterday

The Sky Simulator is a crucial tool to validate the instrument



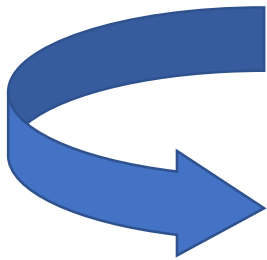
Example of map making with Sky-Simulator for KISS



III. Laboratory tests and characterization

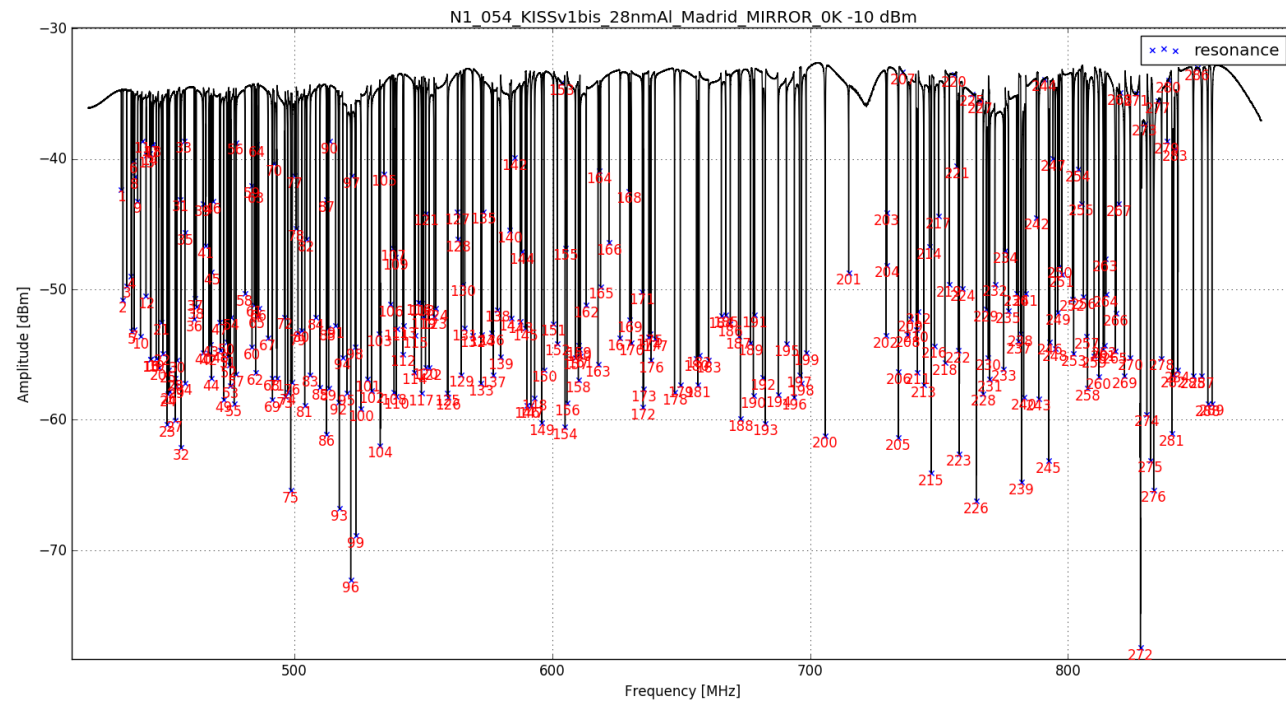
Geometry of the array

Each pixel
is a frequency

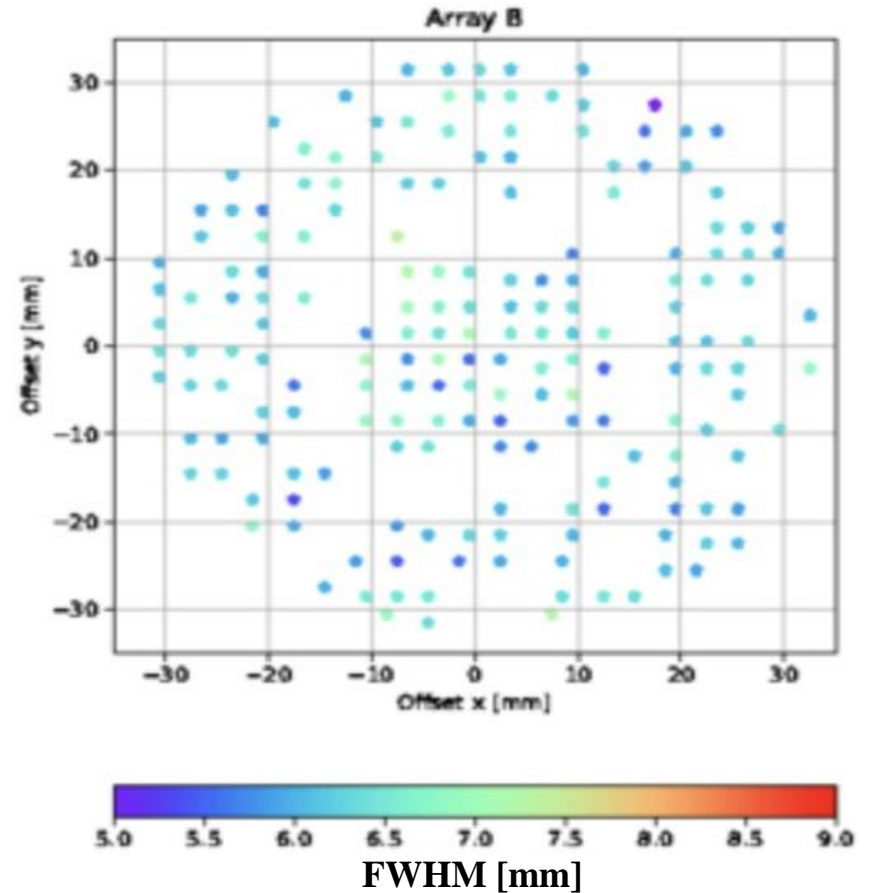


Geometric distribution constrained
with the sky-simulator

23	18	0	35
28	34	6	27
14	1	21	11
33	8	26	16



Geometrical characterization of the array



IV. Installation and observations

November 2018 – NOW

IV. Installation and observations

Chronological path

Installation [November 2018- January 2019]:

- transportation of the instrument (4-days journey)
- mechanical, electrical and network installation
- interface with the telescope

Operation in situ [February 2019 – June 2019]:

- maintenance in situ
- commissioning phase

Operation remotely [June 2019 – ¿October? 2019]:

- maintenance remotely (technicians formed)
- observations [in progress]



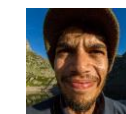
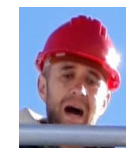
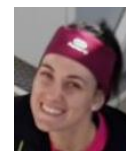
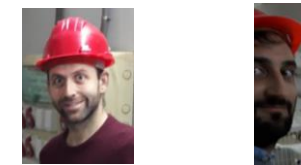
KISS and the Teide vulcano

IV. Installation and observations

A long way to Tenerife



INSTALLATION TEAM



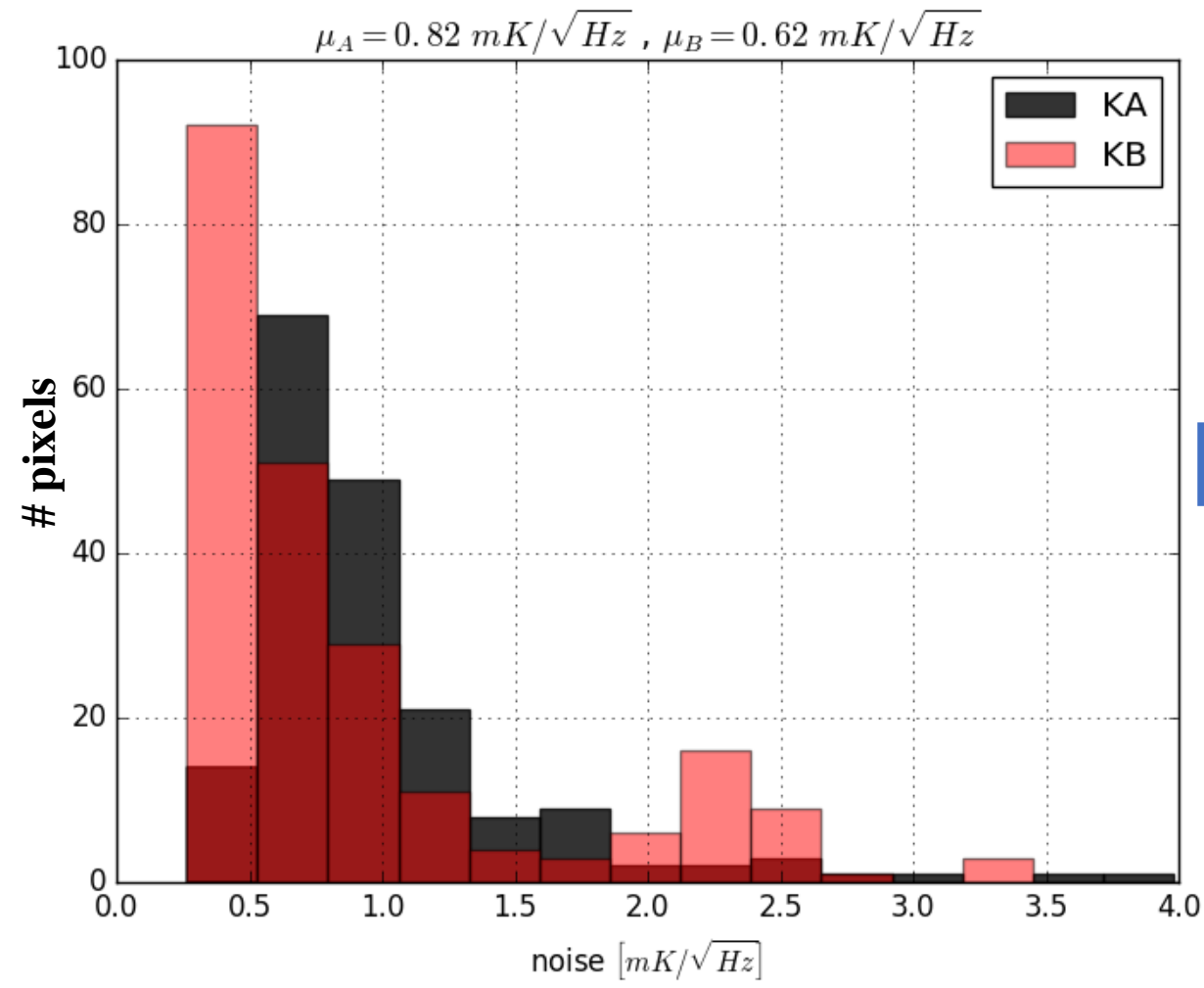
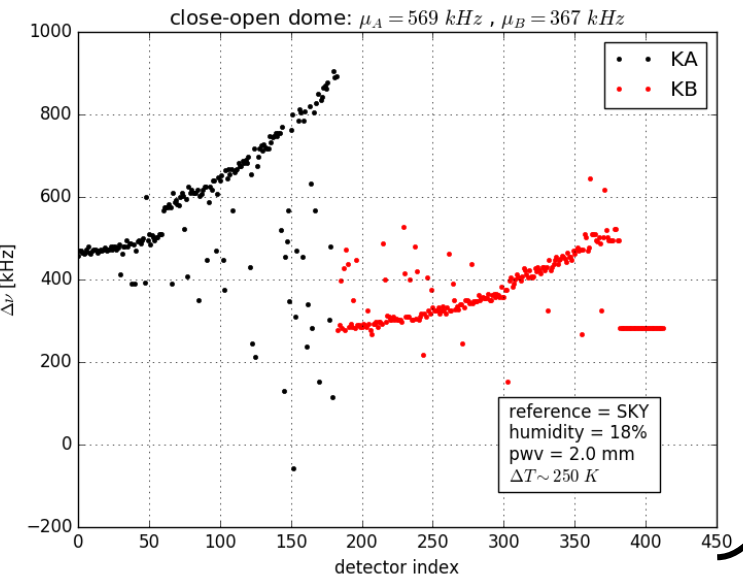
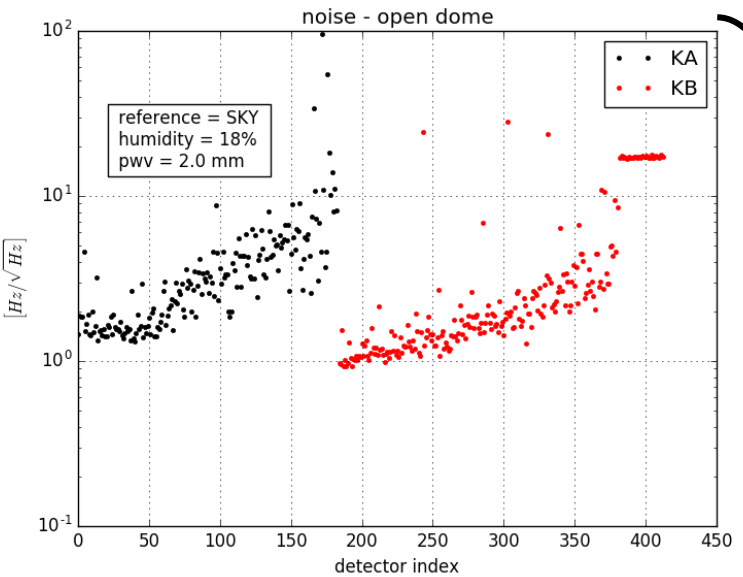
IV. Installation and observations

Sensitivity characterization

PRELIMINARY
TEST

$$S = \Delta\nu / \mathfrak{R}$$

$$\mathfrak{R} \approx 1.5 \text{ kHz/K}$$



IV. Installation and observations

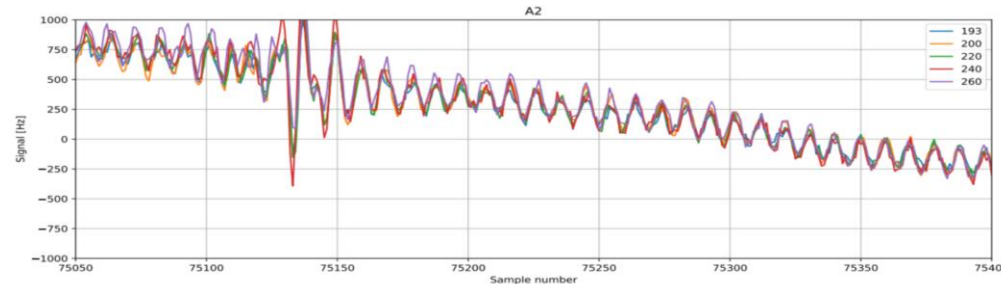
Facts about the observation

FACTS

Commissioning phase in progress
and we are handling three major issues

1) Noise:

source of noise at 400 Hz.



2) Pointing Model:

pointing correction are of the order of degree especially at higher elevation angles. They are due to the different distribution of weight to respect to the previous instrument installed at the same telescope. We do not have many sources to calibrate.

3) Point source detection:

we did not detect a point source up to now. We are investigating the best focus and the possibility that something is wrong with the optics.

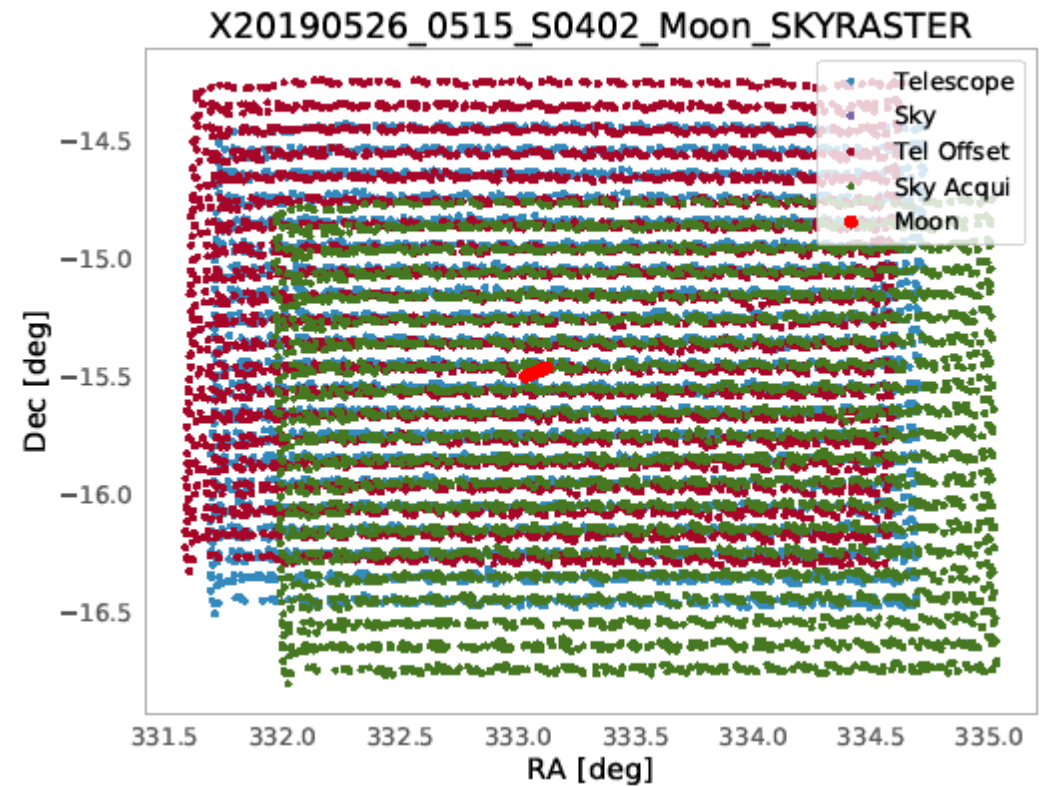
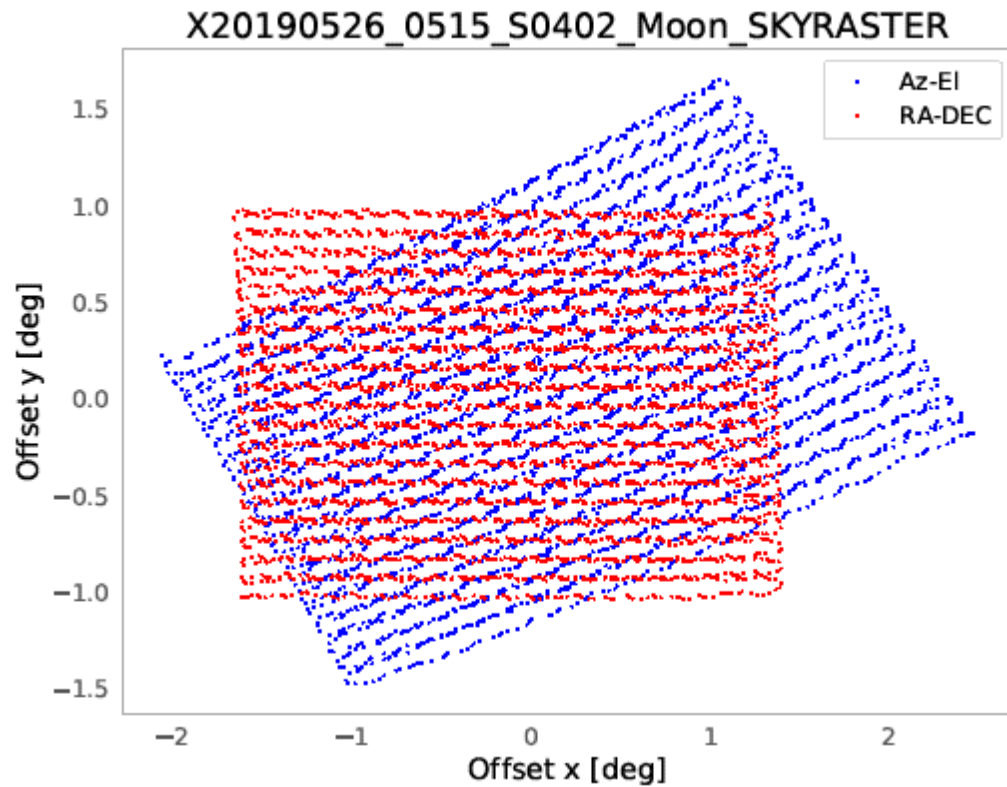
IV. Installation and observations

Moon observation

SIZE = 180 ' x 120 '

STEP = 3 '

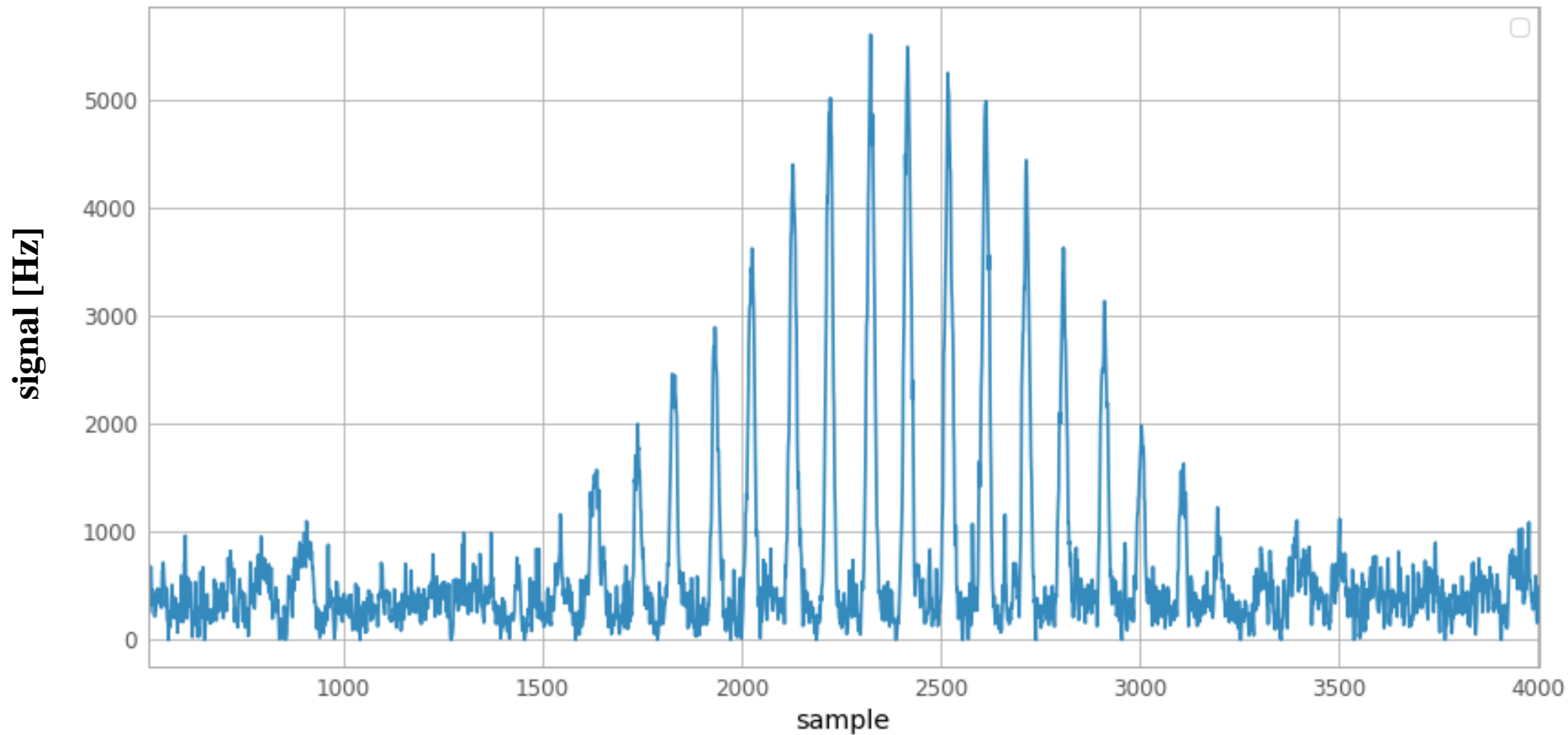
SPEED = 7 ' / s



IV. Installation and observations

Observations of the Moon

Single pixel timeline

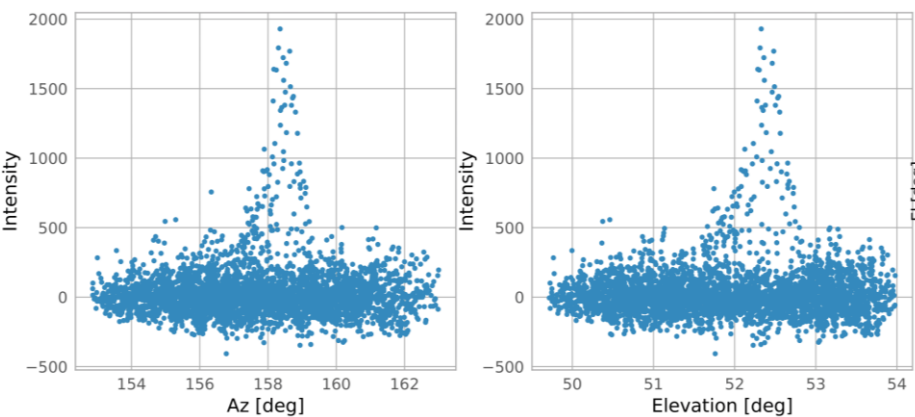


IV. Installation and observations

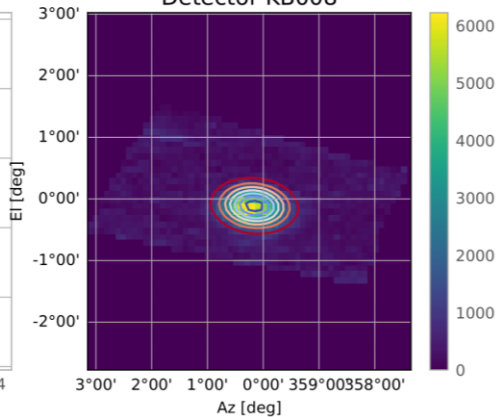
Observations of the Moon

Single pixel map

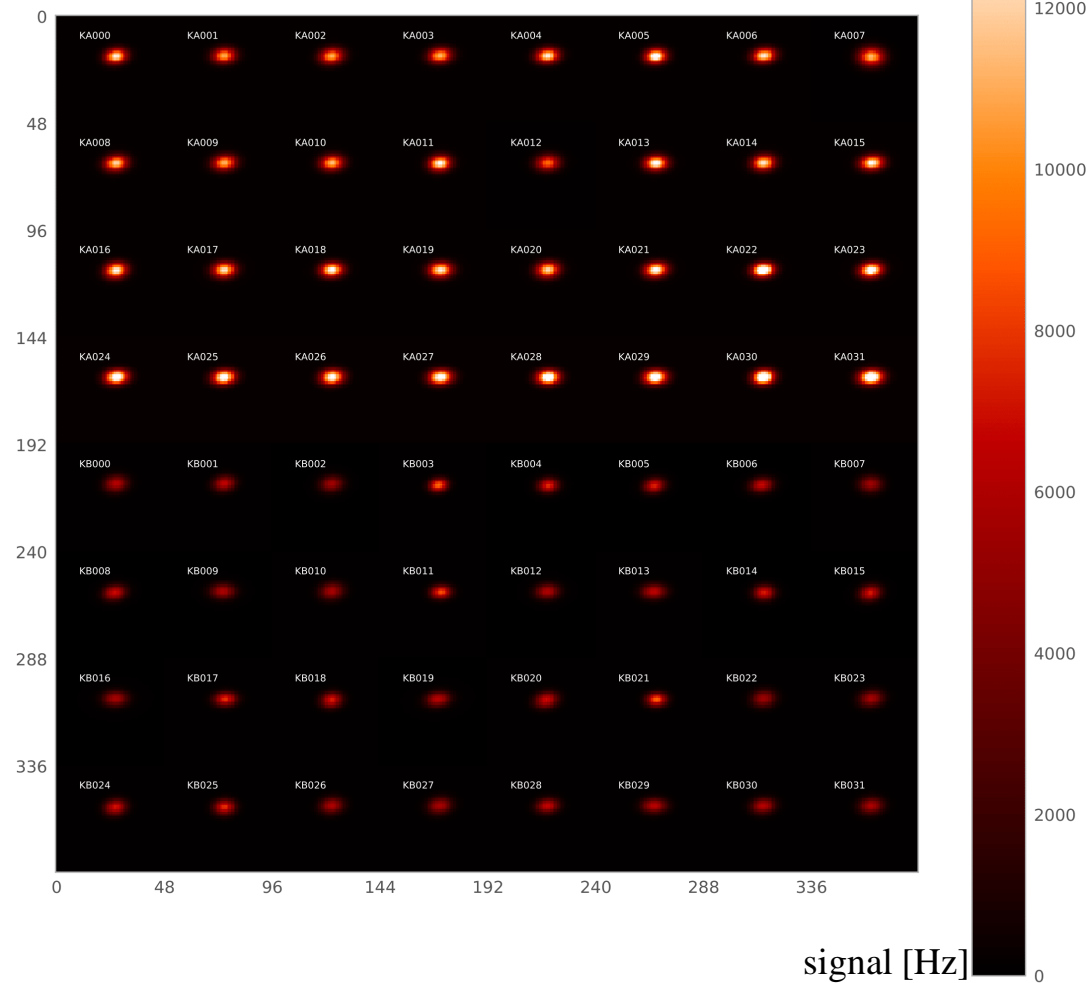
signal [Hz]



Detector KB008



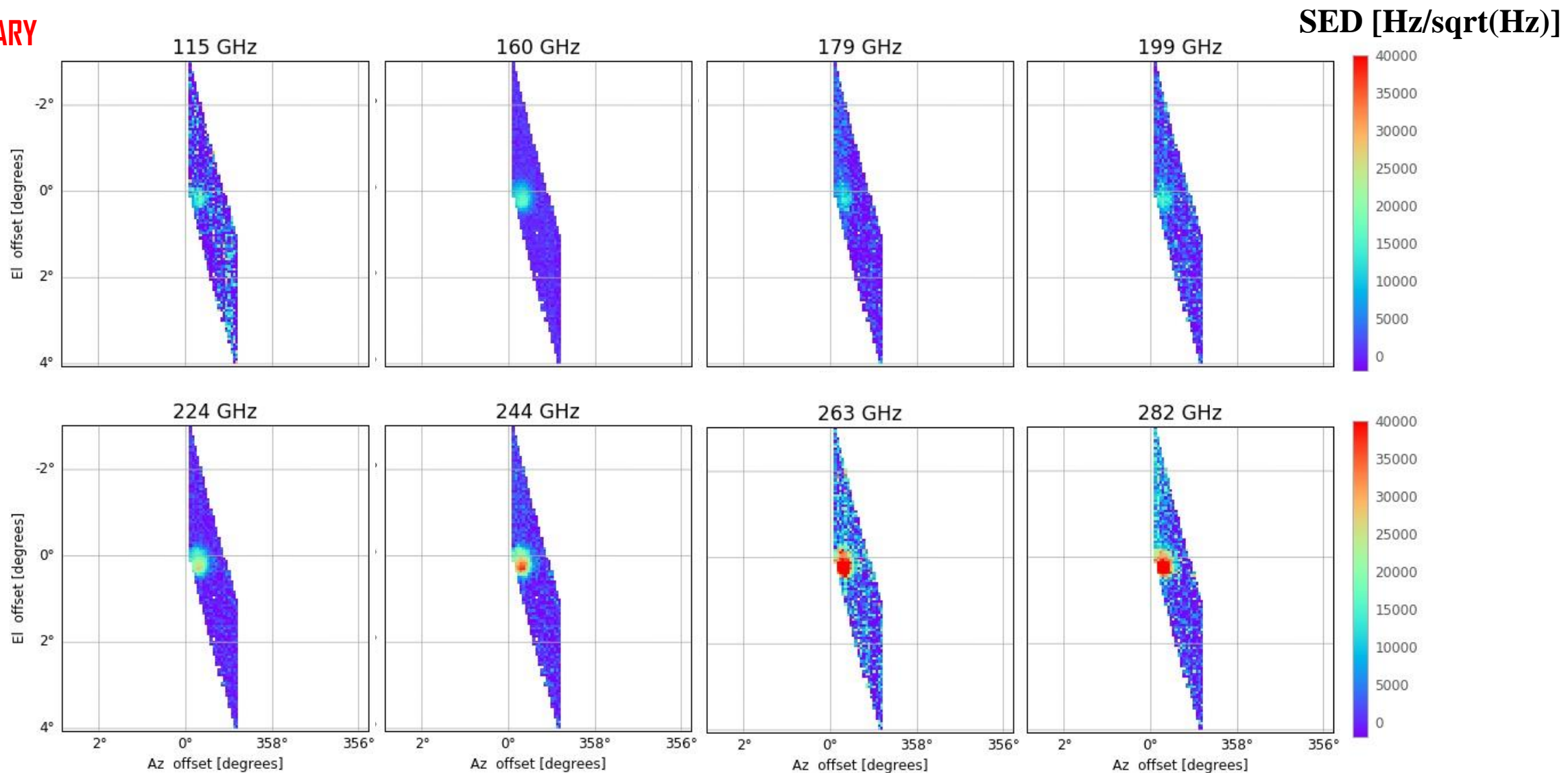
Whole array map



IV. Installation and observations

Moon spectrum-maps

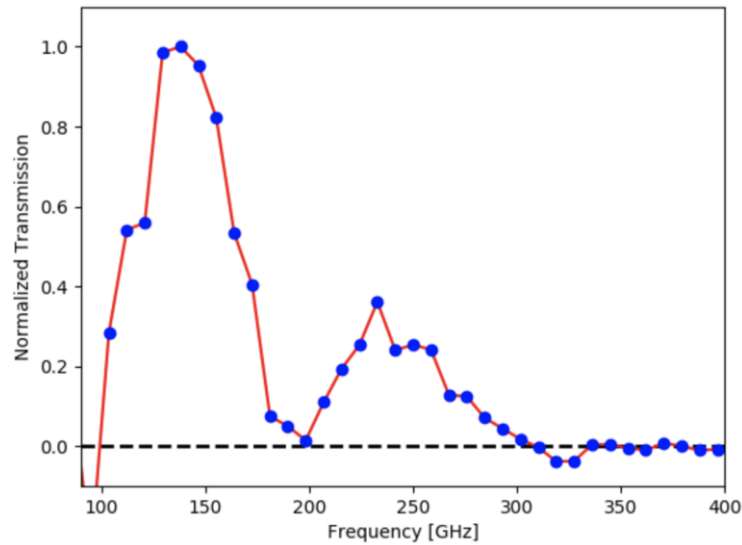
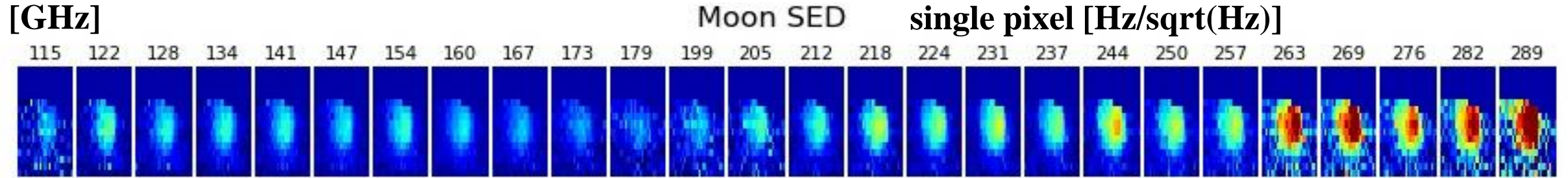
PRELIMINARY



IV. Installation and observations

Moon spectrum-maps

PRELIMINARY



Demonstration of the mapping capability on multi-wavelength

V. Conclusions and perspectives

Past, present and future of the observations

WHAT IS DONE

- installation
- sensitivity on the sky
- demonstration of mapping capability on multi-frequency

WHAT NOW

- point source detection
- calibration on sources (geometry, beam, skydips ...)

WHAT IN THE FUTURE

- full characterisation of the instrument
- observation of clusters of galaxies
- technology and pipeline transfer to CONCERTO
 - same concept of transfer from NIKA to NIKA2
 - [see next talk of G. Lagache]



"Le Penseur" (Auguste Rodin)